

## LOCAL LUMINOUS INFRARED GALAXIES. III. CO-EVOLUTION OF BLACK HOLE GROWTH AND STAR FORMATION ACTIVITY?

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## ABSTRACT

Local luminous infrared (IR) galaxies (LIRGs) have both high star formation rates (SFR) and a high AGN (Seyfert and AGN/starburst composite) incidence. Therefore, they are ideal candidates to explore the co-evolution of black hole (BH) growth and star formation (SF) activity, not necessarily associated with major mergers. Here, we use *Spitzer*/IRS spectroscopy of a complete volume-limited sample of local LIRGs (distances of  $< 78$  Mpc). We estimate typical BH masses of  $3 \times 10^7 M_{\odot}$  using [Ne III]15.56  $\mu$ m and optical [O III] $\lambda$ 5007 gas velocity dispersions and literature stellar velocity dispersions. We find that in a large fraction of local LIRGs the current SFR is taking place not only in the inner nuclear  $\sim 1.5$  kpc region, as estimated from the nuclear 11.3  $\mu$ m PAH luminosities, but also in the host galaxy. We next use the ratios between the SFRs and BH accretion rates (BHAR) to study whether the SF activity and BH growth are contemporaneous in local LIRGs. On average, local LIRGs have SFR to BHAR ratios higher than those of optically selected Seyferts of similar AGN luminosities. However, the majority of the IR-bright galaxies in the RSA Seyfert sample behave like local LIRGs. Moreover, the AGN incidence tends to be higher in local LIRGs with the lowest SFRs. All this suggests that in local LIRGs there is a distinct IR-bright star forming phase taking place prior to the bulk of the current BH growth (i.e., AGN phase). The latter is reflected first as a composite and then as a Seyfert, and later as a non-LIRG optically identified Seyfert nucleus with moderate SF in its host galaxy.

*Subject headings:* galaxies: nuclei — galaxies: Seyfert — infrared: galaxies

## 1. INTRODUCTION

One of the most fundamental relations in extragalactic astronomy is that, at least in the local universe, the masses of supermassive black holes (BH, with masses  $M_{\text{BH}} > 10^6 M_{\odot}$ ) correlate with the stellar mass and velocity dispersion of the bulges of their host galaxies (see e.g. Magorrian et al. 1998; Gebhardt et al. 2000; Marconi & Hunt 2003; Häring & Rix 2004). This seems to imply that bulges and supermassive BH evolve together and regulate each other.

Mergers of gas-rich galaxies are efficient in both producing elevated star formation rates (SFR) and transporting gas to the nuclear region to allow for BH growth. If sufficient matter becomes available very close to the nuclear BH and is accreted, the nucleus of the galaxy will shine with enormous power as an active galactic nucleus (AGN, Lynden-Bell 1969). In parallel, and probably before the fully developed AGN phase, the merger will

trigger a high rate of star formation leading to a luminous infrared (IR)-active phase (Sanders et al. 1988), either as a luminous IR galaxy (LIRG) or ultraluminous IR galaxy (ULIRG). These are defined as having an IR  $8 - 1000 \mu\text{m}$  luminosity  $L_{\text{IR}} > 10^{11} L_{\odot}$  or  $L_{\text{IR}} > 10^{12} L_{\odot}$ , respectively (see Sanders & Mirabel 1996, for a review).

As an alternative to this sequence, Kormendy et al. (2011) proposed that BHs in bulgeless galaxies and in galaxies with pseudobulges grow through phases of *low-level* Seyfert-like activity. Such growth is believed to be driven stochastically by local processes (secular processes), and thus it would not have a global impact on the host galaxy structure (see also Hopkins et al. 2008). Moreover, there is now evidence of two fundamentally different modes of BH growth at work in early- and late-type galaxies in the local universe (Schawinski et al. 2010). The role of moderate luminosity AGN that reside in late-type galaxies and do not involve any recent major mergers is not clear and needs further investigation.

LIRGs are powered by both AGN and star formation activity (Sanders & Mirabel 1996). Alonso-Herrero et al. (2012) showed that while an AGN accompanies star formation activity in a large proportion of local LIRGs, in most cases the AGNs are not energetically important. This is because the AGNs hosted in local LIRGs have Seyfert-like luminosities and thus their bolometric contribution to the total IR luminosity is small (typically 5%). Therefore, the IR luminosities of LIRGs imply SFRs in the range  $11 - 110 M_{\odot} \text{ yr}^{-1}$  using the Kennicutt (1998) prescription converted to a Kroupa (2012) Initial Mass Function (IMF).

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TABLE 1  
SPITZER/IRS CALIBRATION SOURCES

Name	AOR
PNG043.1+37.7	4109056
PNG011.7-00.6	4109568
PNG025.8-17.9	4110080
PNG265.7+04.1	4111104
PNG206.4-40.5	4111616
PNG285.7-14.9	4112128
PNG054.1-12.1	4112640
PNG009.4-05.0	4114176
PNG342.1+10.8	4115200
PNG316.1+08.4	4116224
Calwav-22B-NGC7027	15343104
IRSS-SEPN-N7293-0001	15752960

It is reasonable to assume that the same gas that is used to form stars in the host galaxy can also be used to feed the AGN provided that there is a mechanism able to transport the gas to the inner region (on scales of less than 0.1 pc) of the galaxy (see e.g. the review of Alexander & Hickox 2012). Indeed, it is now apparent that there is a relation between star formation activity on sub-kiloparsec and bulge scales and the BH accretion rate (BHAR), as shown by observations and numerical simulations (Heckman et al. 2004; Hopkins & Quataert 2010; Diamond-Stanic & Rieke 2012). This relation is also predicted to be present, although with a lower significance, for the integrated SFR of the galaxy (Hopkins & Quataert 2010). Therefore, LIRGs show propitious conditions to study the co-evolution of star formation activity and BH growth as they show both high integrated and nuclear SFRs (Alonso-Herrero et al. 2006) and a high occurrence of AGN and composite nuclei (60 – 70%, Yuan et al. 2010; Alonso-Herrero et al. 2012).

In this paper we study the co-evolution of the SFR activity and the BH growth in the complete volume-limited sample of local LIRGs of Alonso-Herrero et al. (2006, 2012). We use *Spitzer* Infrared Spectrograph (IRS, Houck et al. 2004) observations to estimate the nuclear ( $\sim 1.5$  kpc) SFRs and compare them with the integrated values from the IR luminosities. We look for spectroscopically resolved [Ne III]15.56  $\mu$ m lines. Assuming that this line is produced in the narrow line region (NLR) of the AGN, we measure its velocity dispersion and use the Dasyra et al. (2011) relations to obtain the masses of the BH hosted in local LIRGs. We complement these with observations of the optical line [O III] $\lambda$ 5007 and literature values of the stellar velocity dispersion  $\sigma_*$  to obtain further estimates of the BH masses in local LIRGs. All this information allows us to determine the ratios between the SFR (both, nuclear and integrated) and the BHAR in local LIRGs and compare them with those of optically selected Seyfert galaxies. This comparison lets us explore the role of the LIRG-phase in the growth of BH in the local universe. Throughout this paper we assume the following cosmology  $H_0 = 70$  km s $^{-1}$ Mpc $^{-1}$ ,  $\Omega_M = 0.3$  and  $\Omega_\Lambda = 0.7$ .

## 2. SAMPLE, OBSERVATIONS, AND DATA REDUCTION

### 2.1. The sample of local LIRGs

For this work we use the complete volume-limited sample of local LIRGs that was presented and discussed in

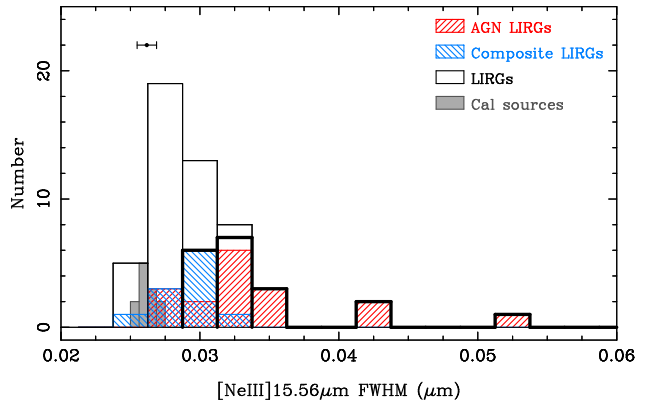


FIG. 1.— Distribution of the observed (not corrected for instrumental resolution) FWHMs of the [Ne III]15.56  $\mu$ m line of local LIRGs compared with those of the calibration sources. The latter are assumed to provide the instrumental resolution at the wavelength of the line. The dot represents the average value of the FWHM of the calibration sources with the corresponding  $1\sigma$  error. We marked those LIRG nuclei classified as AGN (Seyfert and/or [Ne V] emitters) and composites. The rest are H II-like or have no classification. The thick line histogram represents the observed FWHM of those galaxies deemed to have spectroscopically resolved lines.

detail by Alonso-Herrero et al. (2006, 2012). We drew the sample from the *IRAS* Revised Bright Galaxy Sample (RBGS, Sanders et al. 2003) to include all the sources with  $\log(L_{\text{IR}}/L_\odot) \geq 11.05$  and  $v_{\text{hel}} = 2750 - 5300$  km s $^{-1}$ . The sample is composed of 45 *IRAS* systems, with 8 containing multiple nuclei. Therefore the sample includes 53 individual galaxies. For the assumed cosmology the distances are in the range  $\simeq 40 - 78$  Mpc, with a median value of 65 Mpc. The IR luminosities of the individual galaxies<sup>9</sup> are in the range  $\log(L_{\text{IR}}/L_\odot) = 10.64 - 11.67$ , with a median value of  $\log(L_{\text{IR}}/L_\odot) = 11.12$ . All the relevant information for the sample can be found in table 1 of Alonso-Herrero et al. (2012).

Combining a number of optical and mid-infrared (mid-IR) indicators, Alonso-Herrero et al. (2012) derived an AGN detection rate of  $\sim 62\%$  for this complete volume-limited sample of local LIRGs. The derived AGN bolometric luminosities, from the mid-IR spectral decomposition and/or X-ray observations, are in the range  $L_{\text{bol}}(\text{AGN}) = (0.4 - 50) \times 10^{43}$  erg s $^{-1}$  with a median of  $\sim 1.4 \times 10^{43}$  erg s $^{-1}$  (Pereira-Santaella et al. 2011; Alonso-Herrero et al. 2012). However, these AGN are overall only responsible for  $\sim 5\%$  of the IR luminosity emitted by local LIRGs, although the AGN contribution varies from source to source (see Imanishi et al. 2010; Alonso-Herrero et al. 2012). This is in good agreement with Petric et al. (2011) for The Great Observatories All-Sky LIRG Survey (GOALS; Armus et al. 2009).

## 2.2. Spitzer/IRS Observations

### 2.2.1. Observations

We retrieved *Spitzer*/IRS (Houck et al. 2004) spectroscopy of the sample of LIRGs taken with the high

<sup>9</sup> Galaxies significantly below the selection luminosity are members of multiple systems with total luminosities that satisfy the selection criterion.

resolution ( $R \sim 600$ ) short-high (SH) and long-high (LH) modules that cover the  $9.9 - 19.6 \mu\text{m}$  and  $18.7 - 37.2 \mu\text{m}$  spectral ranges, respectively. Details on the programme IDs, the observations, and data reduction of the galaxies, including those from GOALS (Armus et al. 2009), are given in Alonso-Herrero et al. (2012) and Pereira-Santaella et al. (2010a,b). We note that for various reasons two galaxies in this volume-limited sample were not observed with the *Spitzer*/IRS high spectral resolution modules. We extracted the nuclear spectra assuming a point source calibration.

We also retrieved SH and LH spectroscopy from the *Spitzer* archive of 12 calibration sources and planetary nebulae to obtain an accurate estimate of the spectral resolution using fine structure lines (see Section 2.2.3). The names and corresponding Astronomical Observation Request (AOR) numbers are given in Table 1. We reduced the data as for the sample of local LIRGs.

### 2.2.2. Line Fitting

For the fine-structure lines we used Gaussians to measure the line flux, equivalent width (EW), and full width half maximum (FWHM) and a first order polynomial to fit the local continuum (see Pereira-Santaella et al. 2010a, for more details). The fluxes of the  $[\text{Ne II}]12.81 \mu\text{m}$ ,  $[\text{Ne III}]15.56 \mu\text{m}$ , and  $[\text{O IV}]25.89 \mu\text{m}$  lines can be found in Pereira-Santaella et al. (2010b) and Alonso-Herrero et al. (2012). For the  $11.3 \mu\text{m}$  polycyclic aromatic hydrocarbon (PAH) feature first we fitted the local continuum using a linear fit between  $10.6$  and  $11.8 \mu\text{m}$ . After subtracting the continuum, we integrated the flux in the wavelength range between  $10.8$  and  $11.6 \mu\text{m}$ . We list the fluxes, EW, and corresponding errors of the  $11.3 \mu\text{m}$  PAH feature in Table 2.

### 2.2.3. Spectrally resolved $[\text{Ne III}]15.56 \mu\text{m}$ lines

The velocity dispersion of the ionized gas in the NLR allows measuring the BH mass ( $M_{\text{BH}}$ ) in an AGN supplementing reverberation mapping techniques. As discussed by Greene & Ho (2005), the NLR is sufficiently compact to be illuminated by the AGN, while large enough to feel the gravitational potential of the bulge. Therefore, there is a relatively good correlation between the gas and the stellar velocity dispersion around local AGN. In the mid-IR Dasyra et al. (2008, 2011) showed that the velocity dispersion of the fine structure lines  $[\text{S IV}]10.51 \mu\text{m}$ ,  $[\text{Ne III}]15.56 \mu\text{m}$ ,  $[\text{Ne V}]14.32 \mu\text{m}$ , and  $[\text{O IV}]25.89 \mu\text{m}$  of AGN are well correlated with  $M_{\text{BH}}$ .

For our *Spitzer*/IRS spectra we focus on the  $[\text{Ne III}]15.56 \mu\text{m}$  line to look for spectrally resolved line profiles. In AGN the flux and luminosity of this fine-structure line are found to correlate well with those of the mid-IR  $[\text{Ne V}]$  lines at  $14.3$  and  $24.3 \mu\text{m}$  (Gorjian et al. 2007; Pereira-Santaella et al. 2010b). The high ionization potential of the  $[\text{Ne V}]$  lines indicates that they are mostly excited by AGN and thus, the  $[\text{Ne III}]15.56 \mu\text{m}$  emission is likely to be as well. When compared with other mid-IR fine structure lines with relatively high ionization potentials (e.g.,  $[\text{S IV}]10.51 \mu\text{m}$ ,  $[\text{O IV}]25.89 \mu\text{m}$ , and the  $[\text{Ne V}]$  lines), the  $[\text{Ne III}]15.56 \mu\text{m}$  line is a good compromise between ionization potential, critical density<sup>10</sup>, intensity of the line, and a clean region of the

TABLE 2  
FLUXES AND EW OF THE NUCLEAR  $11.3 \mu\text{m}$  PAH FEATURE

Galaxy	Flux	EW
NGC 23	$430.4 \pm 8.5$	$0.78 \pm 0.05$
MCG+12-02-001	$396.8 \pm 7.9$	$0.42 \pm 0.01$
NGC 633	$166.1 \pm 4.0$	$0.61 \pm 0.04$
ESO 297-G012	$128.4 \pm 3.5$	$0.60 \pm 0.05$
UGC 01845	$347.5 \pm 5.2$	$0.74 \pm 0.03$
UGC 02982	$379 \pm 35$	$0.84 \pm 0.15$
CGCG 468-002 NED01	$59.6 \pm 5.6$	$0.24 \pm 0.03$
CGCG 468-002 NED02	$78.4 \pm 7.7$	$0.46 \pm 0.07$
UGC 03351	$275.5 \pm 8.3$	$0.78 \pm 0.06$
NGC 2369	$357.1 \pm 3.9$	$0.58 \pm 0.02$
NGC 2388	$428 \pm 14$	$0.53 \pm 0.04$
MCG+02-20-003	$149.5 \pm 4.1$	$0.70 \pm 0.06$
NGC 3110	$299.0 \pm 4.0$	$0.80 \pm 0.03$
NGC 3256	$1547.2 \pm 4.8$	$0.556 \pm 0.004$
ESO 264-G057	$165 \pm 13$	$0.76 \pm 0.11$
IC 694	$446 \pm 16$	$0.38 \pm 0.02$
NGC 3690	$308 \pm 25$	$0.09 \pm 0.01$
ESO 320-G030	$450.3 \pm 4.6$	$0.78 \pm 0.03$
MCG-02-33-098 W	$99.0 \pm 4.2$	$0.27 \pm 0.02$
MCG-02-33-098 E	$102.3 \pm 6.8$	$0.62 \pm 0.12$
IC 860	$43.2 \pm 2.2$	$0.44 \pm 0.04$
MCG-03-34-064	$23.0 \pm 3.2$	$0.013 \pm 0.002$
NGC 5135	$376.9 \pm 4.0$	$0.46 \pm 0.01$
ESO 173-G015	$434 \pm 31$	$0.53 \pm 0.06$
IC 4280	$249.4 \pm 7.0$	$0.80 \pm 0.04$
UGC 08739	$146 \pm 41$	$0.84 \pm 0.38$
ESO 221-IG010	$259.9 \pm 9.4$	$0.63 \pm 0.05$
NGC 5653	$385 \pm 10$	$0.81 \pm 0.07$
NGC 5734	$402 \pm 11$	$0.80 \pm 0.06$
NGC 5743	$251 \pm 14$	$0.75 \pm 0.12$
IC 4518 E	$64.6 \pm 1.4$	$0.62 \pm 0.05$
IC 4518 W	$62.6 \pm 5.9$	$0.16 \pm 0.02$
Zw 049.057	$62.3 \pm 4.5$	$0.96 \pm 0.23$
NGC 5936	$240.0 \pm 6.6$	$0.67 \pm 0.05$
NGC 5990	$243 \pm 56$	$0.29 \pm 0.09$
NGC 6156	$141 \pm 16$	$0.22 \pm 0.03$
IRAS 17138-1017	$307.6 \pm 4.3$	$0.53 \pm 0.02$
IRAS 17578-0400	$177 \pm 10$	$0.76 \pm 0.09$
IC 4687	$478.2 \pm 4.2$	$0.70 \pm 0.02$
IC 4734	$195.1 \pm 3.8$	$0.59 \pm 0.03$
NGC 6701	$242.2 \pm 4.2$	$0.67 \pm 0.03$
MCG +04-48-002	$394 \pm 33$	$0.73 \pm 0.11$
NGC 7130	$230.5 \pm 3.7$	$0.36 \pm 0.01$
IC 5179	$391.9 \pm 3.6$	$0.68 \pm 0.02$
NGC 7469	$583.5 \pm 7.7$	$0.237 \pm 0.005$
NGC 7591	$142 \pm 12$	$0.36 \pm 0.06$
NGC 7679	$363 \pm 16$	$0.61 \pm 0.05$
NGC 7769	$85 \pm 13$	$0.57 \pm 0.21$
NGC 7770	$131 \pm 18$	$0.47 \pm 0.13$
NGC 7771	$359 \pm 16$	$0.65 \pm 0.07$

**Notes.** The fluxes are in units of  $10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$  and the EW are in units of  $\mu\text{m}$ .

mid-IR spectrum for accurate measurements.

Using the calibration sources (Section 2.2.1) we obtained twelve independent measurements of the unresolved width of the  $[\text{Ne III}]15.56 \mu\text{m}$  line and inferred an instrumental width of  $\text{FWHM}_{\text{inst}} = 0.02618 \pm 0.0007 \mu\text{m}$ . This is equivalent to a spectral resolution of  $R = \lambda/\text{FWHM}_{\text{inst}} \sim 595 \pm 16$  and an instrumental velocity dispersion of  $\sigma_{\text{inst}} = 215 \pm 6 \text{ km s}^{-1}$  at the wave-

ionization potential than the optical  $[\text{O III}]\lambda 5007$  emission line but the optical line has a slightly higher critical density (see figure 6 of Dasyra et al. 2011).

<sup>10</sup> For reference the  $[\text{Ne III}]15.56 \mu\text{m}$  line has a slightly higher

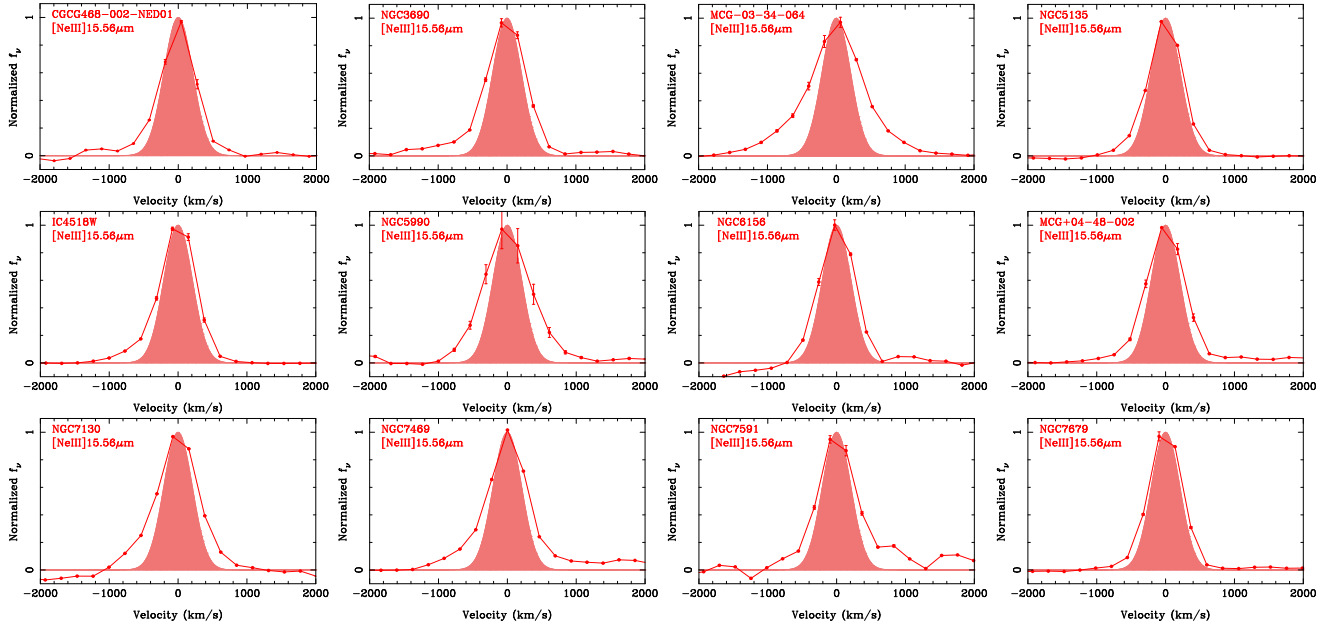


FIG. 2.— Observed profiles (solid lines and filled dots) of the spectrally resolved  $[\text{Ne III}]15.56\ \mu\text{m}$  lines of local LIRGs that are classified as a Seyfert and/or are  $[\text{Ne V}]$  emitters. The shaded area shows a *Spitzer*/IRS SH unresolved profile represented as a Gaussian with  $\text{FWHM} = 0.02618\ \mu\text{m}$ , as determined from *Spitzer*/IRS calibration sources (Section 2.2.1).

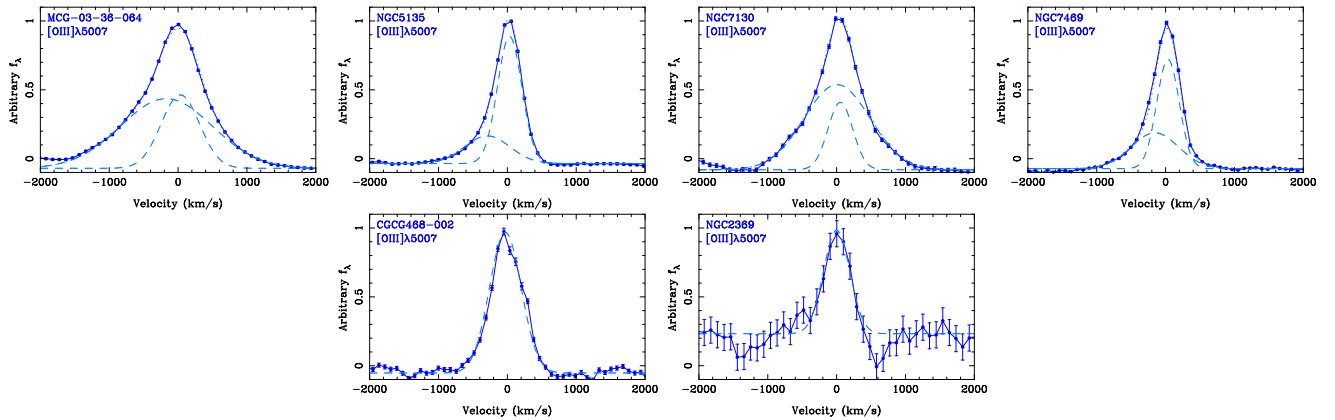


FIG. 3.— **Top panel:** Observed profiles (solid lines and filled dots) of the  $[\text{O III}]\lambda 5007$  line fitted with two Gaussian components (dashed lines) as explained in Section 2.3. **Bottom panel:** Same as top panel but for galaxies whose profiles are fitted with one Gaussian. The plotted errors for the 6dF data (NGC 2369, MCG -03-34-064, NGC 5135, and NGC 7130) are those computed from the rms of the continuum adjacent to the line.

length of this line. While the derived spectral resolution is very similar to that from Dasyra et al. (2011), the standard deviation of our measurements is much lower. Dasyra et al. (2011) obtained  $\sigma_{\text{inst}}$  as the average of the values of the high excitation fine structure lines of a sample of AGN that were deemed not to be resolved by *Spitzer*/IRS with their method. By doing so, they probably also included measurements of barely resolved lines. This would therefore explain the higher dispersion of their instrumental value.

For the LIRGs we need a criterion to assess whether the lines are resolved or not since a significant number of sources have measured widths slightly larger than the instrumental resolution. Figure 1 shows the distribution of the measured FWHM of the  $[\text{Ne III}]15.56\ \mu\text{m}$  line for the sample of LIRGs compared with those of the calibration sources. To avoid analyzing barely resolved lines we adopt the following criterion for the velocity dispersion to determine whether a line is clearly resolved:

$(\sigma_{\text{obs}} - \sigma_{\text{inst}}) > 3 \times \sqrt{\epsilon_{\text{m}}^2 + \epsilon_{\text{inst}}^2}$ , where  $\sigma_{\text{inst}}$  and  $\epsilon_{\text{inst}}$  are the instrumental resolution and its standard deviation, and  $\sigma_{\text{obs}}$  and  $\epsilon_{\text{m}}$  are the measured values for the LIRGs. The total error of the observations  $\epsilon_{\text{obs}}$  includes both the error in the measurement  $\epsilon_{\text{m}}$  and that associated with the instrumental value, that is,  $\epsilon_{\text{obs}} = (\epsilon_{\text{m}}^2 + \epsilon_{\text{inst}}^2)^{1/2}$ .

Using the criterion described above we found that 19 LIRGs in our sample show resolved  $[\text{Ne III}]15.56\ \mu\text{m}$  lines. Most are LIRGs classified as Seyfert (Figure 1) and/or are  $[\text{Ne V}]$  emitters. Those nuclei classified as H II, on the other hand, tend to show FWHM close to the instrumental resolution. Composite nuclei have FWHM between those of unresolved lines and the clearly resolved  $[\text{Ne III}]15.56\ \mu\text{m}$  lines. For those LIRGs deemed to have unresolved  $[\text{Ne III}]15.56\ \mu\text{m}$  lines, the average measured FWHM of the  $[\text{Ne III}]15.56\ \mu\text{m}$  line is  $0.02776 \pm 0.0017\ \mu\text{m}$ . That is, the observed values for these lines are on average  $2\sigma$  above the instrumental resolution.

Table 3 lists the intrinsic values of the velocity

TABLE 3  
VELOCITY DISPERSIONS AND BH MASSES

Galaxy	Class	$\sigma([\text{Ne III}])$ $\text{km s}^{-1}$	$\sigma([\text{O III}])_{\text{n}}$ $\text{km s}^{-1}$	$\sigma([\text{O III}])_{\text{b}}$ $\text{km s}^{-1}$	$\sigma_*$ $\text{km s}^{-1}$	Ref	$\log M_{\text{BH}}$ $M_{\odot}$
Seyfert nuclei and [NeV] emitters							
CGCG 468-002-NED01	Sy2*	$135 \pm 12$	$218 \pm 22$	...	...	...	8.06
NGC 3690	Sy2	$167 \pm 10$	...	...	$144 \pm 11$	1	7.52
MCG -03-34-064	Sy1	$368 \pm 18$	$249 \pm 10$	$656 \pm 26$	155	2	7.65
NGC 5135	Sy2	$138 \pm 9$	$127 \pm 9$	$305 \pm 21$	$124 \pm 6$	3	7.24
IC 4518W	Sy2	$156 \pm 12$	...	...	...	...	7.48
NGC 5990	Sy2	$264 \pm 11$	...	...	...	...	< 8.45
NGC 6156	[Ne v]	$160 \pm 13$	...	...	...	...	7.52
MCG +04-48-002	[Ne v]	$158 \pm 7$	...	...	...	...	7.50
NGC 7130	Sy2:	$259 \pm 16$	$130 \pm 9$	$463 \pm 32$	$147 \pm 5$	3	7.55
NGC 7469	Sy1	$182 \pm 13$	$145 \pm 7$	$336 \pm 17$	$152 \pm 16$	4	7.61
NGC 7591	Sy2:	$169 \pm 17$	...	...	...	...	7.62
NGC 7679	Sy1/Sy2	$102 \pm 7$	176*	...	96	2	6.77
Composite and H II nuclei							
NGC 1614	Composite	$157 \pm 10$	...	...	$164 \pm 8$	5	7.75
NGC 2369	Composite	$133 \pm 8$	$146 \pm 15$	...	...	...	...
IC 694	LINER	< 70	...	...	$141 \pm 17$	5	7.48
NGC 3256	H II	< 112	...	...	$100 \pm 6$	3	6.84
MCG -02-33-098W	Composite	$103 \pm 9$	...	...	...	...	...
ESO 173-G015	...	$102 \pm 7$	...	...	...	...	...
IC 4280	H II	$128 \pm 8$	...	...	...	...	...
IRAS 17138-1017N	Composite	< 80	...	...	$72 \pm 5$	6	6.24
IRAS 17578-0400	...	$119 \pm 8$	...	...	...	...	...
NGC 6701	Composite	< 159	...	...	$150 \pm 20$	7	7.59
NGC 7771	Composite	$178 \pm 17$	...	...	...	...	...

**Notes.**— The references for the spectral class of the nuclei are listed in Alonso-Herrero et al. (2012) except for CGCG 468-002-NED01, which is from the newly analyzed optical spectrum (see Section 2.3).

\*The [O III] velocity dispersion for NGC 7679 is also from reference 2.

References for  $\sigma_*$ : 1. Ho et al. (2009). 2. Gu et al. (2006). 3. Garcia-Rissmann et al. (2005). 4. Onken et al. (2004). 5. Hinz & Rieke (2006). 6. Shier & Fisher (1998). 7. Márquez et al. (1996).

dispersion for the LIRGs with resolved [Ne III]15.56  $\mu\text{m}$  lines as well as their spectroscopic classes from optical spectroscopy and/or mid-IR indicators (Alonso-Herrero et al. 2012). The velocity dispersions (corrected for instrumental resolution) are between 102 and 368  $\text{km s}^{-1}$ . For galaxies in our sample with literature values of the stellar velocity dispersions whose [Ne III]15.56  $\mu\text{m}$  lines are deemed to be unresolved, we provide in this table upper limits to the gas velocity dispersion. In Figure 2 we show the normalized profiles of the twelve Sy/[Ne v] emitter LIRGs with spectrally resolved [Ne III]15.56  $\mu\text{m}$  lines compared with those of an unresolved profile represented as a Gaussian function.

### 2.3. Optical [O III] $\lambda$ 5007 measurements

We obtained archival optical spectra of six LIRGs classified as AGN or Composite from the six-degree Field (6dF) Galaxy Survey (6dFGS, Jones et al. 2004, 2009) and from observations with the FAST spectrograph (Tokarz & Roll 1997; Fabricant et al. 1998) from the Astrophysics (CfA) telescope data center.

NGC 2369, MCG -03-34-06, NGC 5135, and NGC 7130 were observed with the 6dF multi-object fibre (angular diameter of 6''.7) spectrograph on the United Kingdom Schmidt Telescope (UKST). At the typical distance of these LIRGs the spectra cover the central  $\sim 2 \text{ kpc}$ . The instrument spectral resolution around the

[O III] emission lines is  $\sigma_{\text{inst}} \sim 125 \text{ km s}^{-1}$ . Long-slit spectroscopy was available for CGCG 468-002-NED01 and NGC 7469 taken with FAST on the Mt. Hopkins Tillinghast 60-inch telescope as part of different observing programs. The slit width was 3'' and the spectral resolution was  $\sigma_{\text{inst}} \sim 75 \text{ km s}^{-1}$ . None of the spectra are flux calibrated, but they can be used to measure the velocity dispersion of the emission lines as well as their ratios.

We first measured the optical line ratios of CGCG 468-002-NED01 since we did not have an optical classification for this galaxy. We classify this galaxy as a Seyfert 2. In Pereira-Santaella et al. (2010b) and Alonso-Herrero et al. (2012) we detected the mid-IR [Ne v] lines in this galaxy and measured a high [O IV]25.89  $\mu\text{m}$ /[Ne II]12.81  $\mu\text{m}$  line ratio. We used both as evidence of the presence of an AGN in this nucleus. The optical classification confirms this.

For all the galaxies we modeled the [O III] emission lines at 5007Å and 4959Å by fitting a Gaussian profile to each line. The relative position of the Gaussians was fixed according to the rest-frame wavelengths of the two lines. Likewise, the relative intensity was also fixed to the value determined by the atomic parameters. We found that the fit with one component was only satisfactory for CGCG 468-002-NED01 and NGC 2369. For the rest we had to add a broad component to the [O III] lines. The wavelengths and intensities of the narrow (also

called core) component and the broad component were considered independent. To estimate the errors of the velocity dispersions we used the rms of the continuum adjacent to the line for the 6dF spectra and the errors of the spectra for the FAST data. The typical errors of the velocity dispersions of the [O III] lines are less than 10%, and are listed for each galaxy in Table 3.

The top panel of Figure 3 shows the galaxies for which the [O III] $\lambda$ 5007 line fits required two components, whereas the bottom panel shows the profiles of NGC 2369 and CGCG 468-002-NED01. Table 3 lists the velocity dispersion of the core of the line and the broad component, if present, corrected for instrumental resolution. The velocity dispersions of the core components measured in the four galaxies with double components are in good agreement with literature  $\sigma_*$  values, except for MCG -03-34-064.

For the four LIRGs with double components the broad component is blueshifted typically by 50 – 300 km s<sup>-1</sup>. The presence of such asymmetric or broad blue wings in the [O III] $\lambda$ 5007 lines is common in AGN (see e.g., Greene & Ho 2005). The total [O III] emission is dominated by the broad component in MCG-03-34-064 and in NGC 7130, whereas in NGC 5135 and NGC 7469 the core component dominates the emission. In NGC 5135 Bedregal et al. (2009) based on both the relatively broad [Si VI]1.96  $\mu$ m coronal line and its spatial extent suggested the existence of different kinematical components and the possible presence of AGN induced outflows. González Delgado et al. (1998) found blueshifted components in the Ly $\alpha$  and [O III] $\lambda$ 5007 lines of both NGC 7130 and NGC 5135. The velocity dispersion of the broad component of the [O III] $\lambda$ 5007 line of NGC 7130 is similar to the H $\alpha$  broad component ( $\sigma \sim 400$  km s<sup>-1</sup>, Bellocchi et al. 2012). In NGC 7469 Wilson et al. (1986) also detected two components in the [O III] $\lambda$ 5007 line.

### 3. RESULTS

#### 3.1. Black Hole Masses

In Table 3 we list, in addition to the gas velocity dispersions from the [Ne III]15.56  $\mu$ m and optical [O III] $\lambda$ 5007 lines, the stellar velocity dispersion  $\sigma_*$  values from the literature, when available. We also included in this table  $\sigma_*$  values from the literature for other LIRGs in our sample (IC 694, NGC 3256, IRAS 17138–1017, and NGC 6701) whose [Ne III]15.56  $\mu$ m lines appear unresolved at the SH spectral resolution. We note, however, that in the case of on-going major mergers the stellar velocity dispersion vs. BH mass relation might not be applicable.

Dasyra et al. (2008, 2011) showed that mid-IR resolved fine-structure lines can be used to study the NLR of AGN. In particular they demonstrated that the mass of the BH correlates with both the velocity dispersion and the luminosity of the NLR region for a sample of optically selected AGN. The relation is:

$$\log(M_{\text{BH}}/M_{\odot}) = -1.82 + 4.24 \log(\sigma_{\text{NeIII}}/\text{km s}^{-1}) \quad (1)$$

Dasyra et al. (2011) fitted a similar relation for the optical [O III] $\lambda$ 5007 line:

$$\log(M_{\text{BH}}/M_{\odot}) = -1.78 + 4.24 \log(\sigma_{\text{OIII}}/\text{km s}^{-1}) \quad (2)$$

For both relations Dasyra et al. (2011) fixed the slope to the value derived from the correlation between the stellar velocity dispersion and the BH mass by Gültekin et al. (2009):

$$\log(M_{\text{BH}}/M_{\odot}) = 8.12 + 4.24 \log(\sigma_*/200 \text{ km s}^{-1}) \quad (3)$$

The typical uncertainties on the derived  $M_{\text{BH}}$  can be as high as 0.8 dex if using the gas  $\sigma$  based on the rms of the relations (Dasyra et al. 2011) and 0.4 dex if using  $\sigma_*$  (Gültekin et al. 2009).

Before we compute the BH masses, we compare the gas velocity dispersions for the LIRGs with both [Ne III]15.56  $\mu$ m and [O III] $\lambda$ 5007 measurements. As we saw in Section 2.3, in some cases the optical line is fitted with two components. For AGN there is generally a good agreement between  $\sigma_*$  and the gas velocity dispersion from the narrow component of [O III] (Onken et al. 2004; Greene & Ho 2005). Therefore, the core of the [O III] $\lambda$ 5007 line probes the gravitational movements in the NLR and thus can be used to estimate  $M_{\text{BH}}$ . On the other hand, the broad components tend to be blueshifted, may be more influenced by the AGN, and are likely produced by outflows. As can be seen from Table 3, the velocity dispersions from the [Ne III]15.56  $\mu$ m line appear to be intermediate between those of the two optical components. Spectrally resolved blueshifted [Ne III]15.56  $\mu$ m lines (as well as the [Ne V] lines) have also been reported for local ULIRGs (Spoon & Holt 2009), and have been interpreted as the result of outflows. Thus, it is possible that in some cases we might overestimate the BH masses when using the [Ne III]15.56  $\mu$ m line if there are several components. The most suspect galaxy in our sample would be NGC 5990, for which  $\sigma_{\text{NeIII}}$  predicts the most massive BH by far in our sample. We note, however, that the calibration of Dasyra et al. (2011) should take these effects into consideration.

In Table 3 we list  $M_{\text{BH}}$  as calculated from, in this order of preference, the stellar velocity dispersion, the velocity dispersion of the core of the [O III] line, and the [Ne III]15.56  $\mu$ m line velocity dispersion. For those LIRGs with no definitive evidence of AGN activity (i.e., classified as composites, H II or unknown), we only list  $M_{\text{BH}}$  in the second part of the table if we have a value of  $\sigma_*$ . This is because we cannot be sure whether the [Ne III]15.56  $\mu$ m emission comes mostly from the NLR or whether it is produced by star formation (Pereira-Santaella et al. 2010b).

For those LIRGs hosting an actively accreting BH (that is, with clear signs of AGN activity, first part of Table 3) we find typical BH masses of  $3 \times 10^7 M_{\odot}$ , ranging from  $6 \times 10^6 M_{\odot}$  to  $3 \times 10^8 M_{\odot}$ . The estimated BH masses for those LIRGs classified as composite or H II are also in this range. The masses of the LIRG BHs appear to be similar to those of the currently growing BH in the local universe hosted in late-type galaxies (Heckman et al. 2004; Schawinski et al. 2010).

The typical Eddington ratios ( $L_{\text{bol}}(\text{AGN})/L_{\text{Edd}}$ ) for the LIRGs with an estimate of  $M_{\text{BH}}$  are  $2 \times 10^{-2}$ , and range between  $5 \times 10^{-4}$  and 0.1. This means that the

BH in local LIRGs are accreting at a lower efficiency than those in local ULIRG (typical values of 0.08 – 0.4 depending on the method used for estimating the BH mass, Dasyra et al. 2006; Veilleux et al. 2009).

The majority of LIRGs with spectroscopically resolved [Ne III]15.56  $\mu\text{m}$  lines are classified as Seyfert galaxies. However, in local LIRGs there is a high fraction of composite objects (AGN/starburst). Moreover, this fraction remains approximately constant as a function of IR luminosity from galaxies with  $L_{\text{IR}} \sim 10^{10} L_{\odot}$  up to ULIRGs (Yuan et al. 2010). As can be seen from Figure 1, composite nuclei in local LIRGs have values of the FWHM of [Ne III]15.56  $\mu\text{m}$  on average lower than those of Seyfert nuclei but higher than those of nuclei classified as H II and the instrumental resolution. Assuming that these composite nuclei do indeed host an AGN and that the [Ne III]15.56  $\mu\text{m}$  line is probing their NLR, then their BH masses ought to be less than  $\sim 2 - 3 \times 10^7 M_{\odot}$ . The estimated typical AGN bolometric luminosities from the mid-IR spectral decomposition are  $10^{43} \text{ erg s}^{-1}$  or less, so we cannot constrain their Eddington ratios.

The BH of local LIRGs are only marginally less massive than those of local ULIRGs. The BH masses of the latter are in the range  $\sim 10^7 - 5 \times 10^8 M_{\odot}$  (Tacconi et al. 2002; Dasyra et al. 2006; Veilleux et al. 2009). Samples of local ULIRGs are mostly dominated by gas-rich interacting galaxies and major mergers (Sanders & Mirabel 1996), with coalesced ULIRGs having slightly larger BH masses than pre-coalescence ULIRGs (Dasyra et al. 2006). Our sample is both flux and volume-limited and thus is composed mostly of LIRGs with  $L_{\text{IR}} \sim 1 - 2 \times 10^{11} L_{\odot}$  (individual galaxies), which also tend to be spiral galaxies, minor mergers, and galaxies in groups. That is, morphologically local LIRGs are not dominated in numbers by major mergers (see e.g. Sanders & Ishida 2004; Kaviraj 2009; Pereira-Santaella 2012) and their BH masses are similar to those of pre-coalescent ULIRGs (Dasyra et al. 2006).

### 3.2. Star Formation Rates

We computed the SFRs for our sample of LIRGs using a number of IR-based indicators, including the fine-structure [Ne II]12.81  $\mu\text{m}$  line (Roche et al. 1991; Ho & Keto 2007), the 11.3  $\mu\text{m}$  PAH feature (Brandl et al. 2006; Diamond-Stanic & Rieke 2012), and the total 8 – 1000  $\mu\text{m}$  IR luminosity (Kennicutt 1998). Although it is not clear whether the presence of an AGN could destroy the PAH carriers or not, at the typical luminosities of the AGN hosted in local LIRGs the 11.3  $\mu\text{m}$  PAHs do not seem to be affected (Diamond-Stanic & Rieke 2010).

For the integrated SFRs we used  $L_{\text{IR}}$  of the individual galaxies as given in Alonso-Herrero et al. (2012, their table 1) after subtracting the AGN contribution. We used the Kennicutt (1998) relation in terms of the IR luminosity converted to a Kroupa IMF:

$$\text{SFR}(M_{\odot} \text{ yr}^{-1}) = 1.14 \times 10^{-10} (L_{\text{IR}}, L_{\odot}). \quad (4)$$

We obtained total SFRs for the individual galaxies in our sample of between 1 and  $60 M_{\odot} \text{ yr}^{-1}$ .

For the nuclear SFR we used the *Spitzer*/IRS spectroscopy. At the median distance of 65 Mpc for our sample of LIRGs, the IRS SH slit width (4.7 arcsec) sub-

tends typically 1.5 kpc. We use the following recipes put forward by Diamond-Stanic & Rieke (2012) to compute the nuclear SFRs based on the [Ne II]12.81  $\mu\text{m}$  line 11.3  $\mu\text{m}$  PAH feature luminosities:

$$\text{SFR}(M_{\odot} \text{ yr}^{-1}) = 9.6 \times 10^{-9} L(11.3 \mu\text{m}, L_{\odot}) \quad (5)$$

$$\text{SFR}(M_{\odot} \text{ yr}^{-1}) = 8.9 \times 10^{-8} L([\text{Ne II}], L_{\odot}). \quad (6)$$

Diamond-Stanic & Rieke (2012) calibrated Equations 5 and 6 for galaxies with  $L_{\text{IR}} < 10^{11} L_{\odot}$  using the Rieke et al. (2009) templates and a Kroupa IMF. We chose to use these calibrations rather than those for higher IR luminosities because the median value of the IR luminosity of the individual galaxies in our sample is  $1.3 \times 10^{11} L_{\odot}$ . In addition, the width of the IRS SH slit only probes the nuclear regions of our sample of LIRGs and thus lower IR luminosities.

We also note that Equation 5 assumed that the flux of the 11.3  $\mu\text{m}$  PAH feature was measured using PAHFIT (Smith et al. 2007). The spectral coverage ( $\sim 10 - 18 \mu\text{m}$ ) and spectral resolution of the SH data used here are not adequate to use PAHFIT. We measured the PAH fluxes using a local continuum (Section 2 and Table 2). Following Smith et al. (2007) we applied a multiplicative factor of two to the PAH fluxes measured with a local continuum to make a proper comparison with PAHFIT fluxes. We also corrected the 11.3  $\mu\text{m}$  PAH fluxes for extinction using the nuclear strength of the 10  $\mu\text{m}$  silicate feature (from Alonso-Herrero et al. 2012) and the extinction law given in Smith et al. (2007). For most of our local LIRGs this correction is small, as the silicate absorptions are moderate (see e.g. figure 6 in Alonso-Herrero et al. 2012) as compared with the more deeply embedded population of local ULIRGs (Spoon et al. 2007).

While the [Ne II]12.81  $\mu\text{m}$  emission in low luminosity and moderate luminosity AGN is mostly produced by star formation, in high luminosity AGN this emission line can also have an important contribution from the AGN (Pereira-Santaella et al. 2010b). To correct the [Ne II]12.81  $\mu\text{m}$  line flux for any possible AGN contribution, we followed the method put forward by Veilleux et al. (2009). This method reproduces the observed values of the [O IV]25.89  $\mu\text{m}$ /[Ne II]12.81  $\mu\text{m}$  line ratio (given in Alonso-Herrero et al. 2012) as a fractional combination of the typical line ratio of an AGN ( $\sim 4$ ) and that of a starburst galaxy ( $\sim 0.01$ ), where the AGN contribution is the free parameter. Using this method we find that the AGN contribution to the nuclear [Ne II]12.81  $\mu\text{m}$  emission of local LIRGs is small. These corrections are only necessary for 10 LIRGs and the AGN contribution to the nuclear [Ne II]12.81  $\mu\text{m}$  fluxes ranges from 3% to 53% (average  $\sim 18\%$ ).

Figure 4 (left panel) compares the nuclear SFRs as derived from the [Ne II]12.81  $\mu\text{m}$  line (after removing the AGN contribution) and from the 11.3  $\mu\text{m}$  PAH feature. We have color coded symbols by the AGN bolometric contribution to IR luminosity of the galaxies (see Alonso-Herrero et al. 2012). As can be seen from this figure, the majority of LIRGs with small or no ( $< 25\%$ ) AGN contribution show a good agreement between the



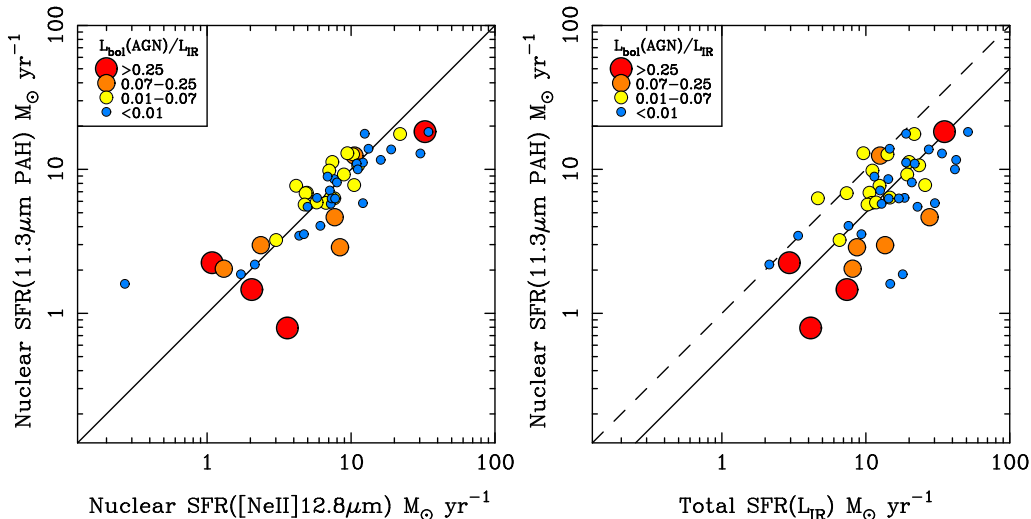


FIG. 4.— Left panel: comparison between the nuclear SFRs from the  $[\text{Ne II}]12.81 \mu\text{m}$  line (corrected for AGN emission) and the  $11.3 \mu\text{m}$  PAH feature for the sample of local LIRGs. The colors denote the AGN bolometric contribution to  $L_{\text{IR}}$  of the individual galaxies:  $< 1\%$  and no AGN contribution (blue),  $1 - 7\%$  (yellow),  $7 - 25\%$  (orange), and  $> 25\%$  (red). The highly discrepant point with no AGN contribution to the left of the straight line is IC 860. The straight solid line is a 1:1 line and is the average value of the ratio between the nuclear SFRs computed with the two estimators (see the text). Right panel: comparison between the nuclear SFR from the  $11.3 \mu\text{m}$  PAH feature and the integrated SFR from the IR luminosity after subtracting the AGN contribution as estimated by Alonso-Herrero et al. (2012). The straight solid line shows an average ratio between the nuclear and the total SFRs of 0.5 (see text), typical of local LIRGs, whereas the dashed line represents a 1:1 relation.

derived nuclear SFRs using the two indicators. Indeed, we find that the ratio between the nuclear SFR estimated from the  $11.3 \mu\text{m}$  PAH feature and that from the  $[\text{Ne II}]12.81 \mu\text{m}$  line is on average  $1.0 \pm 0.3$ .

In a few cases with a significant AGN bolometric contribution the nuclear SFR derived from the  $[\text{Ne II}]12.81 \mu\text{m}$  line is higher than that from the  $11.3 \mu\text{m}$  PAH feature, even after correcting for the AGN contribution. Given the good agreement between the nuclear SFR computed with the two methods for LIRGs with small or no AGN contribution and to avoid AGN contamination issues from now on for the nuclear SFR we use those derived from the  $11.3 \mu\text{m}$  PAH feature. Thus, the individual nuclei of the sample of local LIRGs show values of the nuclear SFR of between  $\sim 0.8$  and  $\sim 20 M_{\odot} \text{ yr}^{-1}$ .

In Figure 4 (right) we compare the total SFR against the nuclear (typically on scales of  $1.5 \text{ kpc}$ ) SFR for the individual galaxies of the LIRG sample. It is clear that a significant fraction of the individual galaxies in our sample have more than half of their total SFRs taking place outside the nuclear ( $1 - 2 \text{ kpc}$ ) regions. For the individual galaxies of the sample the average value of the nuclear SFR over the total SFR ratio is  $0.5 \pm 0.3$ . This agrees with findings using other indicators (Hattori et al. 2004; Alonso-Herrero et al. 2006; Díaz-Santos et al. 2010; Rodríguez-Zaurín et al. 2011). We also note, however, that in approximately 20% of our volume-limited sample of LIRGs the nuclear emission accounts for most of the total SFR measured in the individual galaxies, as is the case for a large fraction of local ULIRGs.

Incidentally, the same comparison for the nuclear SFRs (Figure 4, left) but using the SFR equations calibrated for  $L_{\text{IR}} > 10^{11} L_{\odot}$  (see Diamond-Stanic & Rieke 2012), results in nuclear SFRs from the  $11.3 \mu\text{m}$  PAH feature which are on average a factor of 1.3 higher than those from the  $[\text{Ne II}]12.81 \mu\text{m}$  line. Moreover, the ratio between the nuclear SFR from the  $11.3 \mu\text{m}$  PAH feature

and the total SFR from  $L_{\text{IR}}$  would also be on average 1.3. This gives us confidence that we used the appropriate SFR calibrations for the nuclear rates.

### 3.3. Relation between BH accretion rates and SFR in LIRGs

The presence of an AGN is an unambiguous signpost of a period of BH growth. The AGN luminosity can then be expressed in terms of the BHAR  $\dot{m}_{\text{BH}}$  and the mass-energy conversion efficiency  $\epsilon$  (see Alexander & Hickox 2012, and references therein):

$$\dot{m}_{\text{BH}}(M_{\odot} \text{ yr}^{-1}) = 0.15(0.1/\epsilon)(L_{\text{bol}}/10^{45} \text{ erg s}^{-1}) \quad (7)$$

We used this relation with the typical value of  $\epsilon = 0.1$  in the local universe (Marconi et al. 2004) and  $L_{\text{bol}}(\text{AGN})$  from the mid-IR spectral decomposition (see Alonso-Herrero et al. 2012). We obtained BHAR between  $0.0007$  and  $0.08 M_{\odot} \text{ yr}^{-1}$  for local LIRGs. The uncertainties of the estimates of the BHAR are dominated by the uncertainties of the AGN bolometric luminosities which are typically less than  $0.4 \text{ dex}$  (see Alonso-Herrero et al. 2012, for details).

Having calculated the BHAR we can now compare them with the nuclear and total SFR for the sample of LIRGs. The distribution of  $\text{SFR}/\dot{m}_{\text{BH}}$  can provide clues as to whether AGN and on-going star formation activity are contemporaneous in local LIRGs. Recent numerical simulations by Hopkins (2012) predict a time offset between the peaks of these activities which is thought to depend on the physical scale within the galaxy and the dynamical time of the galaxy, among other parameters.

Figure 5 shows the distribution of  $\text{SFR}/\dot{m}_{\text{BH}}$  for the sample of local LIRGs for the nuclear ( $\sim 1.5 \text{ kpc}$ ) and the integrated SFRs, only for those LIRGs with an estimate of  $L_{\text{bol}}(\text{AGN})$ . The nuclear and integrated  $\text{SFR}/\dot{m}_{\text{BH}}$  are in the range  $\sim 12$  to  $\sim 2 \times 10^4$ , with median values of the nuclear and integrated  $\log(\text{SFR}/\dot{m}_{\text{BH}})$  of 3.1 and 3.4,



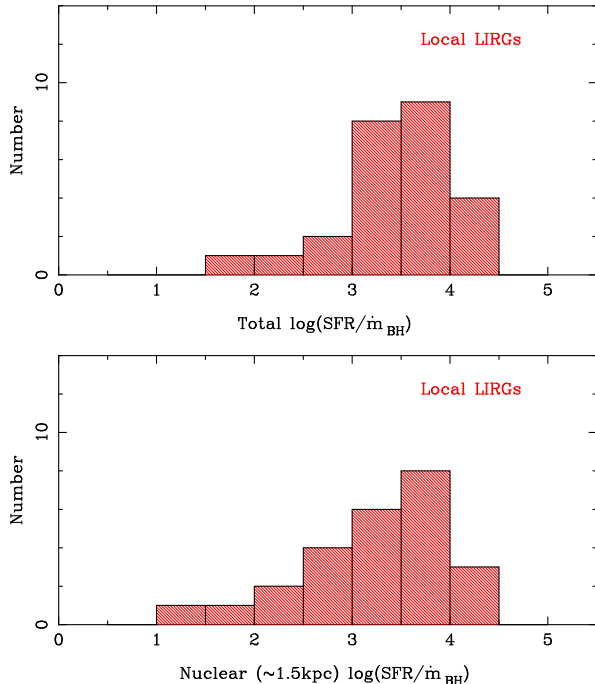


FIG. 5.— Distributions of the  $\text{SFR}/\dot{m}_{\text{BH}}$  ratios for nuclear (typical scales of 1.5 kpc) SFRs (lower panel) and integrated SFRs (upper panel) for the sample of local LIRGs with an estimate of the AGN luminosity. The nuclear SFRs are from the  $11.3\,\mu\text{m}$  PAH feature and the integrated SFR from  $L_{\text{IR}}$  after subtracting the AGN component (see the text for details).

respectively. Most of the nuclei classified as Seyfert from optical spectroscopy show nuclear  $\log(\text{SFR}/\dot{m}_{\text{BH}}) < 3.0$ . The nuclei classified as composite tend to show larger values of the nuclear SFR to BHAR ratio, as expected if they were powered by both SB and AGN activity.

For the entire sample of local LIRGs we obtain  $\sum \text{SFR} / \sum \dot{m}_{\text{BH}} \geq 2900$ . This is a lower limit because for half of the sample we did not detect an AGN component and therefore we are only using an upper limit for their BHAR. If we repeat this only for those LIRGs with an estimate of their AGN bolometric luminosities we obtain a time-averaged value of  $\text{SFR}/\dot{m}_{\text{BH}} \simeq 2500$ , if we assume that all LIRGs go through an AGN phase, that is, a  $\sim 50\%$  duty cycle. The two values are consistent with each other taking into account the typical uncertainties in calculating  $L_{\text{bol}}(\text{AGN})$  ( $\sim 0.4$  dex, Alonso-Herrero et al. 2012). This ratio is a few times higher than that measured for the local population of bulge-dominated galaxies ( $\sim 10^3$ , Heckman et al. 2004) and the local normalization of the BH mass versus bulge mass relation (Marconi & Hunt 2003; Häring & Rix 2004). Even though the SFRs of LIRGs are much higher than those of local bulge galaxies, the time-averaged  $\text{SFR}/\dot{m}_{\text{BH}}$  value of local LIRGs is similar to that of local bulge galaxies during the first 0.3 Gyr of their on-going starburst (Wild et al. 2010, their figure 9). This suggests the existence of similar time delays between the peak of star formation and BH growth in local LIRGs as well.

#### 4. DISCUSSION

##### 4.1. Comparison between the AGNs in LIRGs and the optically identified RSA Seyferts

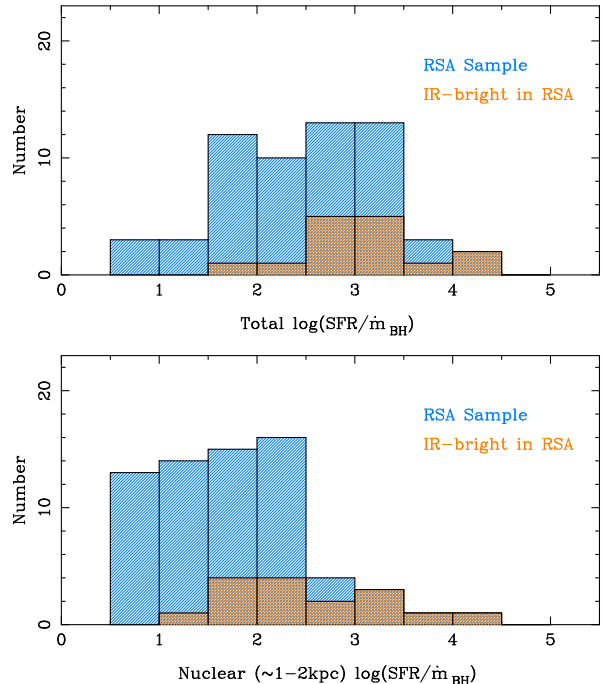


FIG. 6.— Distributions of the ratio between the nuclear SFR and the BHAR (bottom panel) and the integrated SFR and the BHAR (upper panel) for the RSA sample of AGN (blue). In orange we mark those RSA AGN that are also IR-bright galaxies (see text). We only included RSA galaxies with  $\dot{m}_{\text{BH}} > 10^{-4} M_{\odot} \text{yr}^{-1}$ .

In this section we compare the properties of the AGNs identified in local LIRGs with the optically selected Seyfert galaxies in the revised-Shapley-Ames catalog (RSA, Sandage & Tammann 1987; Maiolino & Rieke 1995). This optical sample was spectroscopically selected from the original galaxy-magnitude limited sample of RSA galaxies and is believed not to be biased against low-luminosity Seyfert galaxies. Also, as for the sample of local LIRGs studied in this work, the Seyfert activity in the RSA sample does not appear to be driven primarily by mergers for the majority of the galaxies. Finally, the AGN bolometric luminosities of the RSA Seyferts are similar to those of the local LIRGs.

The median values of  $\text{SFR}/\dot{m}_{\text{BH}}$  for our LIRGs are on average higher than those of local Seyfert galaxies, especially the nuclear values (see Figure 5 and Diamond-Stanic & Rieke 2012). Of course, this is not completely unexpected as local LIRGs, including those hosting an AGN, are essentially selected by their star formation activity, whereas the RSA Seyferts are selected by their host galaxy magnitudes. The latter is more likely probing the stellar mass, rather than the star formation activity.

We can make a more meaningful comparison if we identify the IR-bright Seyferts in the RSA sample. We set the limit to  $\log(L_{\text{IR}}/L_{\odot}) = 10.8$  to match approximately the lower limit covered by the IR luminosities of the individual galaxies of our sample of local LIRGs (see table 1 in Alonso-Herrero et al. 2012). We also included from our sample the newly identified Seyfert nuclei with an estimate of  $L_{\text{bol}}(\text{AGN})$ . The values of  $\text{SFR}/\dot{m}_{\text{BH}}$  for the RSA Seyferts are taken from Diamond-Stanic & Rieke (2012) except for those already included in our sam-

ple of LIRGs. Diamond-Stanic & Rieke (2012) computed the nuclear SFRs from the  $11.3\mu\text{m}$  PAH feature and the MIPS  $24\mu\text{m}$  photometry, and  $L_{\text{bol}}(\text{AGN})$  from the  $[\text{O IV}]\lambda 25.89\mu\text{m}$  line. Finally, for the few RSA Seyferts deemed to have their  $[\text{O IV}]\lambda 25.89\mu\text{m}$  luminosities strongly contaminated by star formation, we estimated  $L_{\text{bol}}(\text{AGN})$  (and thus BHAR) from their hard X-ray luminosities and using a bolometric correction. In this comparison we only included RSA Seyferts with  $\dot{m}_{\text{BH}} > 10^{-4} M_{\odot} \text{yr}^{-1}$ , to match the approximate AGN detection threshold in the LIRG sample. With this limit the IR-bright galaxy fraction in the RSA sample of Seyfert galaxies is then  $\sim 22\% \pm 5\%$ .

Figure 6 shows the distributions of the nuclear and total  $\text{SFR}/\dot{m}_{\text{BH}}$  for the all the RSA Seyferts indicating which ones are also IR-bright. Clearly, those RSA Seyfert classified as IR-bright have ratios similar to, although slightly smaller than, those of the complete sample of local LIRGs. Given the fact that the AGNs in the RSA sample are not very different from those in the LIRGs, the results above would suggest that the bright AGN phase comes after and is somewhat distinct from the LIRG star forming phase. This interpretation would be consistent with observational works showing a delay between the onset of the star formation activity and the latter feeding of the AGN with the consequent BH growth (Davies et al. 2007; Wild et al. 2010) as well as with predictions from numerical simulations (e.g., Hopkins 2012).

#### 4.2. SFRs versus AGN Detection in local LIRGs

It is well known that the AGN fraction increases with increasing IR luminosity from LIRGs to ULIRGs, and at the highest IR luminosities the AGN might dominate bolometrically the luminosity of the system (Veilleux et al. 1995, 2009; Yuan et al. 2010; Nardini et al. 2010; Alonso-Herrero et al. 2012). We now examine further the possibility that the IR-bright AGN phase comes after the distinct IR-bright star-forming phase. If that were the case, we would expect the AGN fraction in LIRGs to be higher at the lowest SFR. The AGN fraction here is computed only for those LIRGs with a secure AGN detection, that is, a Seyfert classification and/or the presence of mid-IR  $[\text{Ne V}]$  lines. As in the previous section, we computed the integrated SFR of the individual galaxies from  $L_{\text{IR}}$  after subtracting the AGN contribution.

Figure 7 shows the AGN fraction as a function of the total SFR of the galaxy. There is a tendency for LIRGs in lowest SFR bin to have a higher AGN incidence ( $36\%^{+14\%}_{-12\%}$ ) than those in the highest SFR bin ( $17\%^{+20\%}_{-11\%}$ ). Alternatively, this could be explained if it is harder to detect AGN in LIRGs with very high SFRs, especially if it takes place mostly in the nuclear regions. Conversely, it may be easier to identify an AGN for low nuclear SFRs and extended star formation. The first explanation is not likely however, as most of the LIRGs in our sample in the high SFR bin also tend to have rather extended SF. In other words, from Figure 4 (right panel) there is no tendency for the fraction of nuclear SFR to increase with the total SFR of the galaxy. In fact, in local LIRGs the mid-IR emission (both from AGN and star formation activity) only starts to be highly concen-

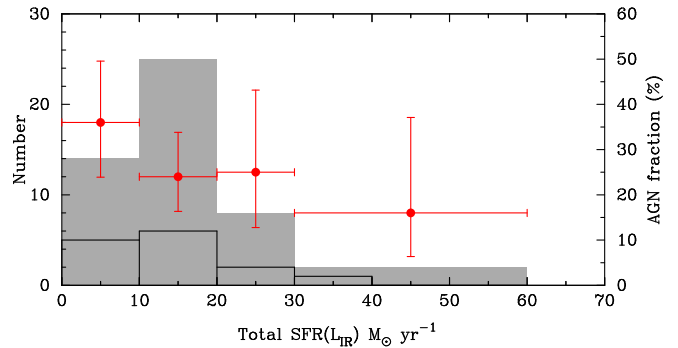


FIG. 7.— Distributions of the integrated SFRs of the local LIRGs (grey histograms) and of those with a Seyfert classification and/or  $[\text{Ne V}]$  emitters (solid line histograms). The data points represent the AGN fraction (scale on the right hand vertical axis) in SFR intervals with  $1\sigma$  error bars for the AGN fraction.

trated in the central regions for IR luminosities above  $10^{11.8} L_{\odot}$ , as shown by Díaz-Santos et al. (2010). We conclude that the higher AGN incidence at low SFRs in local LIRGs provides further evidence for a time delay between the peaks of the star formation and BH growth.

#### 4.3. Key differences between merger and non-merger local LIRGs

Before we explore the role of the LIRG phase in the context of star formation activity and BH growth of massive late-type galaxies (next section), it is important to point out the differences in the star formation histories of merger and non-merger LIRGs in the local universe.

The triggering of the activity in local ULIRGs and likely also in the most luminous local LIRGs is driven by major mergers. As a consequence, ULIRGs tend to host more luminous AGN (quasar or nearly quasar-like luminosities), slightly more massive BH and have higher AGN bolometric contributions than local LIRGs (see Section 3.1 and Veilleux et al. 2009; Nardini et al. 2010; Alonso-Herrero et al. 2012). Local merger LIRGs do show evidence for a recent and intense period of star formation that consumed the gas faster, that is, a *bursty* SFR (see e.g., Alonso-Herrero et al. 2000, 2001) as in local ULIRGs (Rodríguez-Zaurín et al. 2010). In this respect, most of the local merger-LIRGs could be considered as sub-ULIRGs and they might go through a future ULIRG phase or may already have experienced such a phase (Murphy et al. 2001).

Major mergers, however, do not dominate in numbers the population of local LIRGs, especially at  $L_{\text{IR}} < 3 \times 10^{11} L_{\odot}$ , as inferred from morphological studies (see e.g., Sanders & Ishida 2004; Alonso-Herrero et al. 2006; Kaviraj 2009). Therefore, in most local LIRGs the current episodes of star formation activity and BH growth are likely the result of a less *violent* process, e.g., via minor mergers, fly-by companions, and/or secular evolution. Poggianti & Wu (2000) showed that most isolated galaxies in their sample of IR-bright galaxies showed on average more moderate optical Balmer absorption features than the strongly interacting systems in their sample. This suggests that in non-merger LIRGs there is not a strong contribution from a dominant post-starburst stellar population that would be associated with a recent bursty star formation history. This agrees with the finding that most of the mass in non-

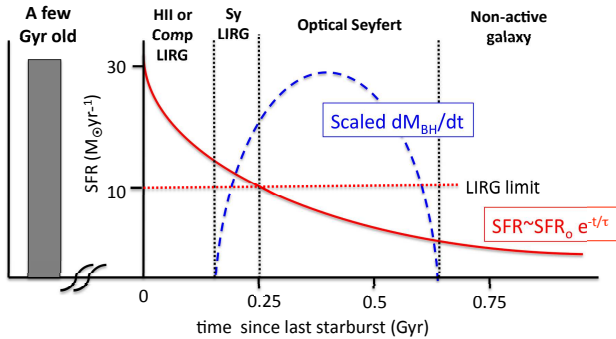


FIG. 8.— Cartoon (not drawn to scale) depicting the proposed scenario the current episode of star formation and BH growth of a local non-major merger LIRG with  $L_{\text{IR}} < 3 \times 10^{11} L_{\odot}$ . The last episode of SF is assumed to be an exponentially decaying SFR (red curve) with an e-folding time of  $\tau = 2.5 \times 10^8$  yr. The current period of BH growth is assumed to have a  $5 \times 10^8$  yr lifetime (blue curve), as in local galaxies with  $M_{\text{BH}} < 10^8 M_{\odot}$  (Marconi et al. 2004). The BHAR would be on average  $\sim 3000$  times lower than the SFR during the IR-bright phase.

merger LIRGs is from evolved stellar populations formed a few Gyr ago with a much smaller contribution in mass from a relatively young stellar population (Kaviraj 2009; Alonso-Herrero et al. 2006, 2010).

#### 4.4. The role of the LIRG-phase in the growth of BH at $L_{\text{IR}} < 3 \times 10^{11} L_{\odot}$

The cartoon in Figure 8 illustrates our proposed scenario for the current episode of BH growth in a non-merger local LIRG typically with  $L_{\text{IR}} < 3 \times 10^{11} L_{\odot}$  (i.e., integrated  $\text{SFR} < 30 M_{\odot} \text{yr}^{-1}$ ). This IR luminosity marks approximately the border between local populations of LIRGs dominated by major merger and populations of disk and relatively isolated LIRGs (Sanders & Ishida 2004).

We assume that the galaxy has an old stellar population formed a few Gyr ago (Kaviraj 2009; Alonso-Herrero et al. 2010) and an existing BH. Since the galaxy did not experience a recent major merger but rather a less *violent* process that induced the last episode of SF, we can assume an exponentially decaying SFR for it, as  $\text{SFR}(t) = \text{SFR}_0 e^{-t/\tau}$  where  $\tau$  is the e-folding time. The peak  $\text{SFR}(t=0) \simeq 30 M_{\odot} \text{yr}^{-1}$  is driven by the  $L_{\text{IR}}$  limit set above. From the position of local AGN on the color-magnitude diagram, Schawinski et al. (2009) demonstrated that the e-folding time cannot be too short (a few Myr or bursty) or too long (i.e., constant SFR). The high AGN detection rate in local LIRGs (Veilleux et al. 1995; Yuan et al. 2010; Alonso-Herrero et al. 2012) implies that LIRGs are able to feed the BH (i.e., show AGN activity) roughly during half of the LIRG-phase (defined as having integrated  $\text{SFR} > 10 M_{\odot} \text{yr}^{-1}$ , see Figure 4). Moreover, given the moderate BH masses inferred for LIRGs (Section 3.1), this AGN activity is likely to be sustained for  $\sim 5 \times 10^8$  yr, that is, the typical lifetime of the BH growth for  $M_{\text{BH}} < 10^8 M_{\odot}$  (see Marconi et al. 2004).

In this simple scenario, the galaxy would be first clas-

sified optically as an H II LIRG and would show elevated integrated and nuclear SFRs. Provided that there is a mechanism to bring the gas near the BH gravitational potential, after a time the existing BH starts accreting material at a sufficiently high rate. Then the galaxy would still have a Seyfert-like AGN luminosity. Because there is an elevated nuclear SFR, the galaxy would be optically classified first as composite (AGN/SB) and later on as a Seyfert LIRG as the SFR decreases. In both cases the  $\text{SFR}/\dot{m}_{\text{BH}}$  ratios would be on average higher than those of optically identified Seyferts. After a few hundred million years the SFR would be below  $10 M_{\odot} \text{yr}^{-1}$  and thus it would not be classified as a LIRG any longer. However, the AGN phase may be still active for another few hundred million years as Seyfert-like AGNs do not require large amounts of gas to remain active (Hopkins 2012). At this point the galaxy would be identified as an optical Seyfert with a lower value of the SFR/BHAR ratio and a moderate SFR. Eventually both the SF and the BHAR would be low and the galaxy may not be identified as a Seyfert. A SFR e-folding time of  $\tau \sim 2.5 \times 10^8$  yr and a BH growth lifetime of  $\sim 5 \times 10^8$  yr would reproduce the SFRs as well as the Seyfert and composite fractions of local LIRGs (Yuan et al. 2010).

In summary, in the local universe LIRGs with  $L_{\text{IR}} < 3 \times 10^{11} L_{\odot}$  LIRGs and an AGN represent an early phase of the (possibly episodic) growth of BHs in massive spiral galaxies with high SFR, not necessarily associated with a major merger event. An H II-like LIRG phase would predate the AGN classified LIRG phase. The latter in turn would be followed by an optically identified AGN phase where the galaxy still shows an appreciable SFR but below the LIRG limit. Although the triggering mechanisms for the SB and AGN activity of LIRGs and ULIRGs might be different, it appears that the sequence of events is similar. Namely, ULIRGs are believed to involve the interaction between two gas-rich galaxies that results first in a high SFR and then an optical QSO phase (see Sanders et al. 1988; Sanders & Mirabel 1996; Yuan et al. 2010, and references therein). The end product, in contrast, is rather different with ULIRGs probably evolving into moderate-mass ellipticals (Sanders & Mirabel 1996; Dasyra et al. 2006; Hopkins et al. 2008) and non-merger LIRGs likely remaining late-type galaxies.

#### 5. SUMMARY

We have studied the current co-evolution of the star formation activity and BH growth for a complete sample of 51 local LIRGs (see Alonso-Herrero et al. 2006, 2012, for details on the sample). We used *Spitzer*/IRS SH ( $R \sim 600$ ) data in the  $\sim 10 - 18 \mu\text{m}$  spectral range to estimate BH masses from resolved [Ne III]  $15.56 \mu\text{m}$  fine structure lines, and to measure nuclear SFRs from the  $11.3 \mu\text{m}$  PAH feature and the [Ne II]  $12.81 \mu\text{m}$  line.

We detected spectrally resolved [Ne III]  $15.56 \mu\text{m}$  lines in 19 LIRGs with gas velocity dispersions between 102 and  $368 \text{ km s}^{-1}$ . Most of these are LIRGs classified as Seyferts and/or are [Ne V] emitters. We assumed that [Ne III]  $15.56 \mu\text{m}$  probes gas arising from the NLR in the AGN LIRGs and then used the Dasyra et al. (2011) relation between  $M_{\text{BH}}$  and [Ne III]  $15.56 \mu\text{m}$  velocity dispersion. We combined gas velocity dispersions from [Ne III]  $15.56 \mu\text{m}$  and [O III]  $\lambda 5007 \text{ \AA}$  and stellar velocity

dispersions from the literature to obtain BH masses between  $M_{\text{BH}} = 6 \times 10^6 M_{\odot}$  and  $M_{\text{BH}} = 3.5 \times 10^8 M_{\odot}$  (median value of  $3 \times 10^7 M_{\odot}$ ). The derived AGN bolometric luminosities of  $L_{\text{bol}}(\text{AGN}) = (0.4-50) \times 10^{43} \text{ erg s}^{-1}$  (median of  $\sim 1.4 \times 10^{43} \text{ erg s}^{-1}$ , see Alonso-Herrero et al. 2012) imply BHAR of  $\dot{m}_{\text{BH}} \simeq 0.0007-0.08 M_{\odot} \text{ yr}^{-1}$  and relatively low Eddington ratios (typically  $2 \times 10^{-2}$ ).

Although AGNs accompany the star formation activity in the large majority of local LIRGs (Veilleux et al. 1995; Yuan et al. 2010; Alonso-Herrero et al. 2012), in most cases the AGN is not energetically dominant (Alonso-Herrero et al. 2012). Using the  $11.3 \mu\text{m}$  PAH feature we obtained nuclear (on typical scales of  $1.5 \text{ kpc}$ ) SFRs for the individual galaxies of between  $\sim 0.8$  and  $\sim 20 M_{\odot} \text{ yr}^{-1}$  and using  $L_{\text{IR}}$  integrated SFRs between  $1$  and  $60 M_{\odot} \text{ yr}^{-1}$  (for a Kroupa IMF). In a large fraction of local LIRGs this on-going star formation activity is not only taking place in the nuclear regions but it is also spread out on kpc scales throughout the galaxy. From the IR estimates in this work the median value for the sample of local LIRGs is  $\text{SFR}(\text{nuclear} \sim 1.5 \text{ kpc})/\text{SFR}(\text{total}) = 0.5$ .

We found that local LIRGs have a time-averaged value of  $\text{SFR}/\dot{m}_{\text{BH}} \geq 2900$ , which is a few times higher than that of the local population of bulge-dominated galaxies ( $\sim 10^3$ , Heckman et al. 2004) and the local normalization of the BH mass versus bulge mass relation (Marconi & Hunt 2003; Häring & Rix 2004). However, this ratio is similar to the value of local bulge galaxies averaged during the first  $\sim 0.3 \text{ Gyr}$  of the last starburst (Wild et al. 2010), that is, before the peak of the BH growth.

Local LIRGs have on average higher  $\text{SFR}/\dot{m}_{\text{BH}}$  ratios (both nuclear and total SFR) than the optically selected RSA Seyferts in the local Universe. However, the IR-bright RSA Seyferts have similar ratios to those of local LIRGs. Since the overall properties of the AGNs in local LIRGs and in the RSA sample are very similar, we concluded that the AGN phase comes after and is somewhat distinct from the LIRG star forming phase. The AGN

fraction in local LIRGs appear to be a function of the integrated SFR, with LIRGs in the low SFR bin having a higher AGN (i.e., Seyfert) incidence than those with high SFRs. This provides further support to the scenario that in local LIRGs the peak of intense IR-bright SF activity occurs prior to the AGN phase.

Finally we put forward a simple model to explore the role of the LIRG-phase in the last episode of BH growth at  $L_{\text{IR}} < 3 \times 10^{11} L_{\odot}$ . At these luminosities most LIRGs are not major-mergers, and have an old stellar population formed a few Gyr ago and a current episode of star formation. The latter only a small mass fraction of the total stellar mass. We propose that an exponentially decaying SFR with a peak  $\text{SFR}(t=0) \simeq 30 M_{\odot} \text{ yr}^{-1}$  and an e-folding time of  $\tau \sim 2.5 \times 10^8 \text{ yr}$  together with a  $5 \times 10^8 \text{ yr}$  period of BH growth would explain the observed SFRs and fractions of Seyfert and composite objects in local LIRGs. In this scenario non-merger LIRGs hosting an AGN represent an early phase of the possibly episodic growth of the BH which is followed by an optically Seyfert phase in massive galaxies with relatively high SFR in their host galaxies.

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#### REFERENCES

- Alexander, D. M., & Hickox, R. C. 2012, *NewAR*, 56, 93  
 Alonso-Herrero, A., Pereira-Santaella, M., Rieke, G. H., & Rigopoulou, D. 2012, *ApJ*, 744, 2  
 Alonso-Herrero, A., García-Marín, M., Rodríguez-Zaurín, J., Monreal-Ibero, A., Colina, L., & Arribas, S. 2010, *A&A*, 522, A7  
 Alonso-Herrero, A., Rieke, G. H., Rieke, M. J., Colina, L., Pérez-González, P. G., & Ryder, S. D. 2006, *ApJ*, 650, 835  
 Alonso-Herrero, A., Engelbracht, C. W., Rieke, M. J., Rieke, G. H., & Quillen, A. C. 2001, *ApJ*, 546, 952  
 Alonso-Herrero, A., Rieke, G. H., Rieke, M. J., & Scoville, N. Z. 2000, *ApJ*, 532, 845  
 Armus, L., Mazzarella, J. M., Evans, A. S., et al. 2009, *PASP*, 121, 559  
 Bedregal, A. G., Colina, L., Alonso-Herrero, A., & Arribas, S. 2009, *ApJ*, 698, 1852  
 Bellocchi, E., Arribas, S., & Colina, L. 2012, *A&A*, 542, 54  
 Brandl, B., Bernard-Salas, J., Spoon, H. W. W., et al. 2006, *ApJ*, 653, 1129  
 Dasyra, K. M., Ho, L. C., Netzer, H., Combes, F., Trakhtenbrot, B., Strum, E., Armus, L., & Elbaz, D. 2011, *ApJ*, 740, 94  
 Dasyra, K. M., Ho, L. C., Armus, L., et al. 2008, *ApJ*, 674, L9  
 Dasyra, K. M., Tacconi, L. J., Davies, R. I., et al. 2006, *ApJ*, 651, 835  
 Davies, R. I., Müller Sánchez, F., Genzel, R., Tacconi, L. J., Hicks, E. K. S., Friedrich, S., & Sternberg, A. 2007, *ApJ*, 671, 1388  
 Diamond-Stanic, A. M., & Rieke, G. H. 2010, *ApJ*, 724, 140  
 Diamond-Stanic, A. M., & Rieke, G. H. 2012, *ApJ*, 746, 168  
 Diamond-Stanic, A. M., Rieke, G. H., & Rigby, J. R. 2009, *ApJ*, 698, 623  
 Díaz-Santos, T., Charmandaris, V., Armus, L., et al. 2010, *ApJ*, 723, 993  
 Fabricant, D., Cheimets, P., Caldwell, N., & Geary, J. 1998, *PASP*, 110, 79  
 García-Rissmann, A., Vega, L. R., Asari, N. V., et al. 2005, *MNRAS*, 359, 765  
 Gebhardt, K., Kormendy, J., Ho, L. C., et al. 2000, *ApJ*, 543, L5  
 González Delgado, R. M. et al. 1998, *ApJ*, 505, 174  
 Gorjian, V., Cleary, K., Werner, M. W., & Lawrence, C. R. 2007, *ApJ*, 655, L73  
 Greene, J. E., & Ho, L. C. 2005, *ApJ*, 627, 721  
 Gu, Q., Melnick, J., Cid Fernandes, R., Kunth, D., Terlevich, E., & Terlevich, R. 2006, *MNRAS*, 366, 480  
 Gültekin, K., Richstone, D. O., Gebhardt, K., et al. 2009, *ApJ*, 698, 198  
 Häring, N., & Rix, H. W. 2004, *ApJ*, 604, L89  
 Hattori, T., Yoshida, M., Ohtani, H., et al. 2004, *AJ*, 127, 736

- Heckman, T. M., Kauffmann, G., Brinchmann, J., Charlot, S., Tremonti, C., & White, S. D. M. 2004, *ApJ*, 613, 109
- Hinz, J. L., & Rieke, G. H. 2006, *ApJ*, 646, 872
- Ho, L. & Keto, E. 2007, *ApJ*, 658, 314
- Ho, L. C., Green, J. E., Filippenko, A. V., & Sargent, W. L. W. 2009, *ApJS*, 183, 1
- Hopkins, P. F. 2012, *MNRAS*, 420, L8
- Hopkins, P. F., & Quataert, E. 2010, *MNRAS*, 407, 1529
- Hopkins, P. F., Hernquist, L., Cox, T. J., & Keres, D. 2008, *ApJS*, 175, 356
- Houck, J. R., Roellig, T. L., van Cleve, J., et al. 2004, *ApJS*, 154, 18
- Imanishi, M., Nakagawa, T., Shirahata, M., Ohya, Y., & Onaka, T. 2010, *ApJ*, 721, 1233
- Jones, D. H., Read, M. A., Saunders, W., et al. 2009, *MNRAS*, 399, 683
- Jones, D. H., Saunders, W., Colless, M., et al. 2004, *MNRAS*, 355, 747
- Kaviraj, S. 2009, *MNRAS*, 394, 1167
- Kennicutt, R. C. Jr. 1998, *ARA&A*, 36, 189
- Kormendy, J., Bender, R., & Cornell, M. E. 2011, *Nature*, 469, 374
- Kroupa, P. 2002, *Science*, 295, 82
- Lynden-Bell, D. 1969, *Nature*, 223, 690
- Magorrian, J., Tremaine, S., Richstone, D., et al. 1998, *AJ*, 115, 2285
- Maiolino, R. & Rieke, G. H. 1995, *ApJ*, 454, 95
- Marconi, A., & Hunt, L. K. 2003, *ApJ*, 589, L21
- Marconi, A., Risaliti, G., Gilli, R., Hunt, L. K., Maiolino, R., & Salvati, M. 2004, *MNRAS*, 351, 169
- Márquez, I., Moles, M., & Masegosa, J. 1996, *A&A*, 310, 401
- Murphy, T. W., Soifer, B. T., Matthews, K., & Armus, L. 2001, *ApJ*, 559, 201
- Nardini, E., Risaliti, G., Watabe, Y., Salvati, M., & Sani, E. 2010, *MNRAS*, 405, 2505
- Onken, C. A., Ferrarese, L., Merritt, D. et al. 2004, *ApJ*, 615, 645
- Pereira-Santaella, M., Alonso-Herrero, A., Rieke, G. H., Colina, L. et al. 2010a, *ApJS*, 188, 447
- Pereira-Santaella, M., Diamond-Stanic, A. M., Alonso-Herrero, A., & Rieke, G. H. 2010b, *ApJ*, 725, 2270
- Pereira-Santaella, M., Alonso-Herrero, A., Santos-Lleo, M. et al. 2011, *A&A*, 535, A93
- Pereira-Santaella, M. 2012, PhD Thesis, Universidad Autónoma de Madrid, Spain
- Petric, A. O., Armus, L., Howell, J., et al. 2011, *ApJ*, 730, 28
- Poggianti, B. M., & Wu, H. 2000, *ApJ*, 529, 157
- Rieke, G. H., Alonso-Herrero, A., Werner, B. J., et al. 2009, *ApJ*, 692, 556
- Roche, P. F., Aitken, D. K., Smith, C. H., & Ward, M. J. 1991, *MNRAS*, 248, 606
- Rodríguez-Zaurín, J., Tadhunter, C., & González Delgado, R. M. 2010, *MNRAS*, 403, 1317
- Rodríguez-Zaurín, J., Arribas, S., Monreal-Ibero, A., Colina, L., Alonso-Herrero, A., & Alfonso-Garzón, J. 2011, *A&A*, 527, 60
- Sandage, A. & Tammann, G. A. 1987, *A Revised Shapley-Ames Catalog of Bright Galaxies*, 2nd ed. (Washington, DC: Carnegie Institution of Washington)
- Sanders, D., & Ishida C., 2004, in Aalto S., Huttemeister S., Pedlar A., eds, *The Neutral ISM in Starburst Galaxies Vol. 320 of Astronomical Society of the Pacific Conference Series*. pp 230
- Sanders, D. B., Mazzarella, J. M., Kim, D.-C., Surace, J. A., & Soifer, B. T. 2003, *AJ*, 126, 1607
- Sanders, D. B., & Mirabel, F. I. 1996, *ARA&A*, 34, 749
- Sanders, D. B., Soifer, B. T., Elias, J. H., Madore, B. F., Matthews, K., Neugebauer, G., & Scoville, N. Z. 1988, *ApJ*, 325, 74
- Schawinski, K., Urry, M. C., Virani, S., et al. 2010, *ApJ*, 711, 284
- Schawinski, K., Virani, S., Simmons, V., Urry, M. C., Treister, E., Kaviraj, S., & Kushkuley, B. 2009, *ApJ*, 692, L19
- Shier, L. M., & Fisher, J. 1998, *ApJ*, 497, 163
- Smith, J. D. T., Draine, B. T., Dale, D. A., et al. 2007, *ApJ*, 656, 770
- Spoon, H. W. W., & Holt, J. 2009, *ApJ*, 702, L42
- Spoon, H. W. W., Marshall, J. A., Houck, J. R., et al. 2007, *ApJ*, 654, L49
- Tacconi, L. J., Genzel, R., Lutz, D., Rigopoulou, D., Baker, A. J., Iserlohe, C., & Tecza, M. 2002, *ApJ*, 580, 73
- Tokarz, S., & Roll, J. 1997, *ASPC*, 125, 140
- Veilleux, S., Kim, D.-C., Sanders, D. B., Mazzarella, J. M., & Soifer, B. T. 1995, *ApJS*, 98, 171
- Veilleux, S., Rupke, D. S. N., Kim, D.-C., et al. 2009, *ApJS*, 182, 628
- Wild, V., Heckman, T., & Charlot, S. 2010, *MNRAS*, 405, 933
- Wilson, A. S., Baldwin, J. A., Sun, S.-D., & Wright, A. E. 1986, *ApJ*, 310, 121
- Yuan, T.-T., Kewley, L. J., & Sanders, D. B. 2010, *ApJ*, 709, 884